

## The APACHE Project

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**Abstract.** First, we summarize the four-year long efforts undertaken to build the final setup of the APACHE Project, a photometric transit search for small-size planets orbiting bright, low-mass M dwarfs. Next, we describe the present status of the APACHE survey, officially started in July 2012 at the site of the Astronomical Observatory of the Autonomous Region of the Aosta Valley, in the Western Italian Alps. Finally, we briefly discuss the potentially far-reaching consequences of a multi-technique characterization program of the (potentially planet-bearing) APACHE targets.

## 1 Introduction

M dwarf stars, with masses  $M_{\star} \leq 0.6 M_{\odot}$ , make up the vast majority of the reservoir of nearby stars within  $\sim 25\text{--}30$  pc. These stars have not traditionally been included in large numbers in the target lists of radial-velocity (RV) searches for planets for two main reasons: 1) their intrinsic faintness, which prevented Doppler surveys in the optical from achieving very high radial-velocity precision ( $< 5\text{--}10$  m/s) for large samples of M dwarfs (e.g., Eggenberger & Udry 2010, and references therein), and 2) their being considered as providers of very inhospitable environments for potentially habitable planets (e.g., Tarter et al. 2007; Scalo et al. 2007, and references therein). These two paradigms are now shifting. There are several important reasons for such a change in perspective, which can be summarized under two main themes.

First, the observational evidence gathered by ultra-high-precision space-borne photometric surveys (e.g., Kepler) indicates that the frequency of low-mass planets, i.e. Neptunes and Super-Earths, is an increasing function of decreasing stellar mass (Howard et al. 2012; Dressing & Charbonneau 2013). This result has recently been strengthened by the findings of ground-based RV programs carried out with state-of-the-art facilities (e.g., HARPS): Super-Earths with  $M_p \leq 10 M_{\oplus}$  within the Habitable Zone<sup>1</sup> (HZ) of low-mass stars appear ubiquitous (Bonfils et al. 2013). Given that the identification of a rocky, habitable planet is the essential prerequisite to its possible characterization as an

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<sup>1</sup>In its standard definition, the Habitable Zone corresponds to the range of distances from a given star for which water could be found in liquid form on a planetary surface (Kasting et al. 1993)

actual life-bearing celestial object, it is thus clear why low-mass M dwarfs, seen for long as providers of inhospitable environments for life (Scalo et al. 2007, and references therein), are now being moved at the center of the stage in the exoplanets arena. Consensus is growing among the astronomers' community that the first habitable rocky planet will be discovered (and might have been discovered already!) around a red M dwarf in the backyard of our Solar System.

Second, the sample of the nearest ( $d < 25 - 30$  pc), relatively bright ( $J < 9 - 10$ ) M dwarfs is amenable to combined studies with a wide array of observational techniques, which can be exploited at the best of their potential providing the opportunity to characterize the architecture of planetary systems across orders of magnitude in mass and orbital separations in a way that's not readily achievable for Solar analogs. For example, the possibility to reach detection of short-period transiting rocky planets from the ground with modest-size telescopes (30–50 cm class) is guaranteed by the small radii of M dwarfs, leading to deep transits ( $\Delta mag \gtrsim 0.005$  mag) for the case of planets with  $2 \leq R_p \leq 4 R_\oplus$  (e.g., Charbonneau et al. 2009). In addition, as we have discussed above, the favorable mass ratios allow for detection of rocky, potentially habitable planets with the RV technique, thanks to the moderately large amplitudes of the RV signals (a few m/s). Analogously, at intermediate separations ( $\approx 1 - 4$  AU) high-precision astrometry becomes sensitive to planets in the mass range between Neptune and Jupiter (e.g., Casertano et al. 2008). Finally, the favorable planet-star contrast ratios provided by the low intrinsic luminosity of M dwarfs allows for improved detectability thresholds of giant planets at wide separations ( $> 5 - 10$  AU) with direct imaging techniques (e.g., Bowler et al. 2012). For the same reason, atmospheric characterization of transiting low-mass Super-Earths can be achieved for this sample (e.g., Tessenyi et al. 2012). Systematic studies of planetary systems orbiting low-mass stars can thus crucially inform planet formation and evolution, structural and atmospheric models, particularly when seen in connection with the properties of the hosts (e.g., Ida & Lin 2005; Mordasini et al. 2012a,b; Seager & Deming 2010, and references therein).

However, not all physical properties of low-mass stars are known precisely enough for the purpose of the detection and characterization of small-radius planets. Worse still, some of the characteristics intrinsic to late-type dwarfs can constitute a significant source of confusion in the interpretation in planet detection and characterization measurements across a range of techniques. First of all, there exist discrepancies between theory and observations in the determination of the sizes of M dwarfs, typically on the order of 10%-15% (Ribas 2006; Torres et al. 2010, and references therein). Second, there are at present difficulties in spectroscopically determining with a high degree of precision M dwarf metallicities, as M-dwarf spectra are dominated by chemically complex molecular features that make the identification of the continuum in an M dwarf spectrum challenging, rendering line-based metallicity indicators unreliable (Woolf et al. 2009; Rojas-Ayala et al. 2010). Third, our understanding of the age-rotation-activity connection for field M dwarfs with age greater than  $t \sim 0.5$  Gyr (e.g., Jenkins et al. 2009; West & Basri 2009) is still subject to rather large uncertainties due to the sparseness of the data. Fourth, as measurements of chromospheric activity indicators ( $H\alpha$  line) have shown how the fraction of active M dwarfs increases as a function of spectral sub-type (e.g., Bochanski et al. 2005; West et al. 2011), activity-related phenomena such as stellar spots, plagues, and flares become increasingly a matter of concern for planet detection and characterization programs targeting late-type stars.

All the above considerations clearly underline how achieving the goal of the detection *and* characterization of low-mass, potentially habitable, rocky planets around low-mass stars requires the construction of a large (all-sky) sample of nearby, relatively bright M dwarfs with well-characterized properties. This will necessitate the combined use of time-series of spectroscopic, astrometric, and photometric data of high quality. In particular, the jitter levels will have to be quantified in detail for each target individually, as the jitter properties may vary from star to star within the same spectral



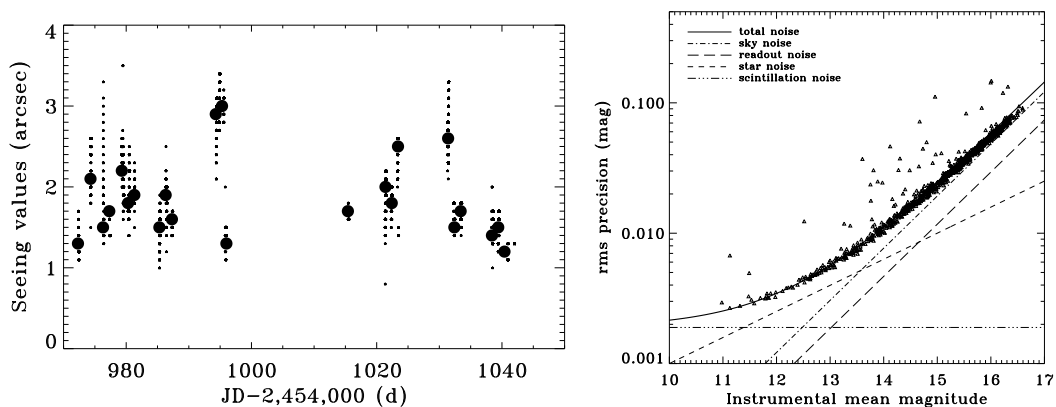
**Figure 1.** Left: Map of northern Italy with the red cross marking the approximate location of the OAVdA in the Western Italian Alps. Right: A night view of the Aosta Valley Observatory.

class, as suggested by recent findings based on high-precision Kepler photometry (e.g., Ciardi et al. 2011) and high-resolution, high-S/N spectroscopy (e.g., Zechmeister et al. 2009). In the following, we summarize the preparations for and present first results from APACHE (A PATHway to the Characterization of Habitable Earths), a new ground-based photometric transiting search for small-size planets around thousands of bright, nearby early- to mid-M dwarfs. The APACHE project is a collaboration between INAF-Osservatorio Astrofisico di Torino (INAF-OATo) and the Astronomical Observatory of the Autonomous Region of the Aosta Valley (OAVdA). APACHE utilizes an array of five 40-cm telescopes at the OAVdA site, in the western Italian Alps (Fig. 1, left and right panels). The survey, with official kick-off in July 2012, will last for five years.

## 2 The Site Characterization Study

In this work (Damasso et al. 2010), we carried out a detailed site characterization study at the OAVdA site, aimed at establishing its potential to host the APACHE project. For the purpose of the site testing campaign, we gathered *R*-band photometric and seeing measurements utilizing different instruments available at the site. We gauged site-dependent observing conditions such as night-sky brightness, photometric precision, and seeing properties. Public meteorological data were also used in order to help in the determination of the actual number of useful observing nights per year.

We measured *V*-band night-sky brightness and extinction levels not dissimilar from those registered at other good observing sites. The median seeing over the period of in situ observations (a few months) was found to be  $\approx 1.7''$  (Fig. 2, left panel), and rather stable. However, given the limited duration of the observations, we did not probe any possible seeing seasonal patterns or the details of its possible dependence on other meteorological parameters, such as wind speed and direction. The fraction of fully clear nights per year amounts to 39%, while the total of useful nights increases to 57%, assuming a (conservative) cloud cover of not more than 50% of the night. Based on the analysis of photometric data collected over the period of 2009 May–August for three stellar fields centered on the transiting planet hosts WASP-3, HAT-P-7, and Gliese 436, we achieved seeing-independent best-case photometric precision  $\sim 3$  mmag (rms) in several nights for bright ( $R < 11$  mag) stars (see Fig. 2, right panel for an example). A median performance  $\sim 6$  mmag during the observing period was obtained for stars with  $R < 13$  mag.

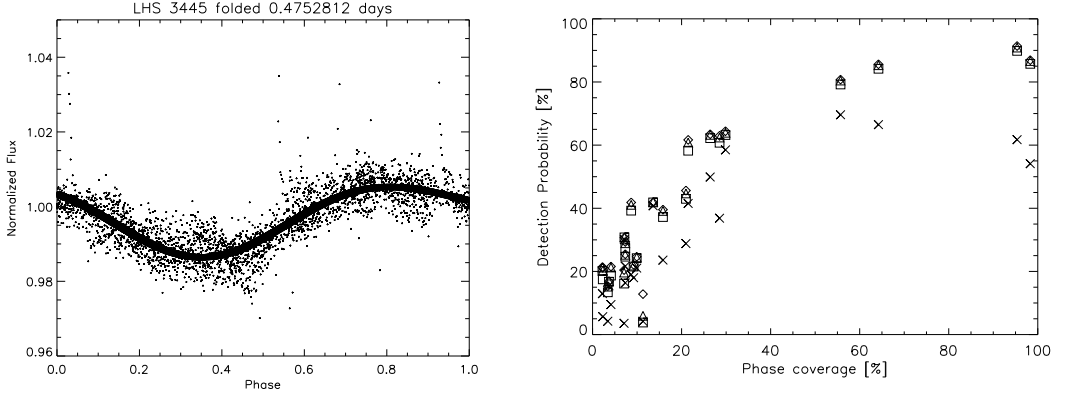


**Figure 2.** Left: The distribution of seeing measurements over the whole period of observation reported in Damasso et al. (2010). The small black stars show individual data points, each corresponding to an average seeing value over a 1 min interval. The large filled circles indicate the median seeing value for each night. Right: Photometric errors (rms) vs. instrumental mean magnitude in the field of the planet-hosting star WASP-3 during a good observing night (Damasso et al. 2010). The various contributions to the expected photometric noise, according to an analytic error budget model, are also shown. Pictures reprinted with permission from Damasso et al. (2010).

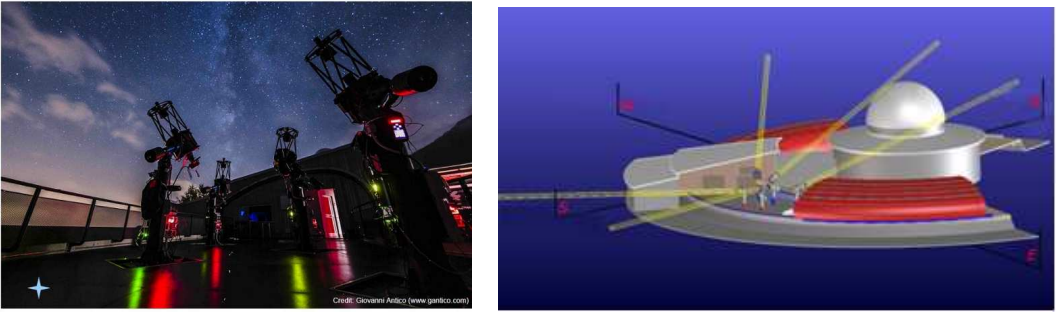
### 3 The Pilot study

The pilot study (Giacobbe et al. 2012) was focused on a year-long photometric monitoring campaign, using small-size instrumentation (25-80 cm class telescope systems), of a sample of 23 bright ( $J < 12$  mag), cool M0-M6 dwarfs with good parallaxes from the TORINO Parallax Program (TOPP; Smart et al. 2010). The primary objectives we set out to achieve in this study were *a*) to demonstrate sensitivity to  $< 4 R_{\oplus}$  (i.e., smaller than radius of Neptune) transiting planets with periods of a few days around our sample, through a two-fold approach that combines a characterization of the statistical noise properties of our photometry with the determination of transit detection probabilities via simulations, and *b*) where possible, to better our knowledge of some astrophysical properties (e.g., activity, rotation, age) of our targets through a combination of spectroscopic and astrometric information and our differential photometric measurements.

We achieved a typical nightly rms photometric precision of  $\sim 5$  mmag, with little or no dependence on the instrumentation used or on the details of the adopted methods for differential photometry, and gauged the degradation levels in the precision with respect to a pure white noise regime due to the presence of correlated (red) noise in our data. Based on a detailed stellar variability analysis, *a*) we detected no transit-like events (an expected result given the sample size); *b*) we determined  $< 1$  day photometric rotation periods for two of the M dwarfs in the sample, in agreement with the large projected rotational velocities inferred for both stars based on the analysis of archival spectra, and found consistency of the estimated inclinations of the stellar rotation axes with those derived using a simple spot model (see Fig. 3, left panel). For two stars in the sample we recorded short-term, low-amplitude flaring events. Finally, based on simulations of transit signals of given period and amplitude injected in the actual (nightly reduced) photometric data for our sample, we derived a relationship between transit detection probability and phase coverage (see Fig. 3, right panel). We found that, using the BLS search algorithm, even when phase coverage approached 100%, there was a limit to the detection probability of  $\approx 90\%$ . Around program stars with phase coverage  $> 50\%$  we



**Figure 3.** Left: Photometric data for the dM star LHS 3445 folded according to its probable rotation period ( $P = 0.472$  days). Overplotted is a best-fit single-spot model, as discussed in Giacobbe et al. (2012). Right: Average detection probability as a function of average phase coverage for  $P < 1.0$  days. Pictures reprinted with permission from Giacobbe et al. (2012).

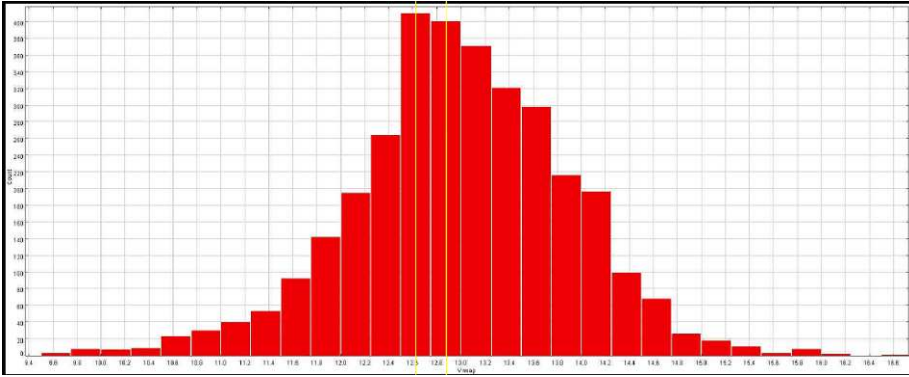


**Figure 4.** Left: Close-up night view of the APACHE telescope array. Right: A cartoon visualizing the exact arrangement of the APACHE telescopes within the OAVdA platform.

would have had  $> 80\%$  chances of detecting planets with  $P < 1$  day inducing fractional transit depths  $> 0.5\%$ , corresponding to minimum detectable radii in the range  $\sim 1.0 - 2.2 R_{\oplus}$ .

## 4 The APACHE Infrastructure

APACHE employs an array of five 40-cm class telescopes hosted on a single platform with electronically controlled roll-off enclosure (see Fig. 4, left panel). The availability of a single movable enclosure and of much of the associated support infrastructure makes the OAVdA a site essentially unique in Europe and will afford a very substantial cost saving for the project. The arrangement of the array has been defined so as to maximize sky coverage during campaign operations (see Fig. 4, right panel). The telescope array is composed of five identical Carbon Truss 40-cm f/8.4 Ritchey-Chrétien telescopes, with a GM2000 10-MICRON mount and equipped with a FLI Proline PL1001E-2 CCD Camera and Johnson-Cousins R & I filters. These systems guarantee state-of-the-art quality perfor-



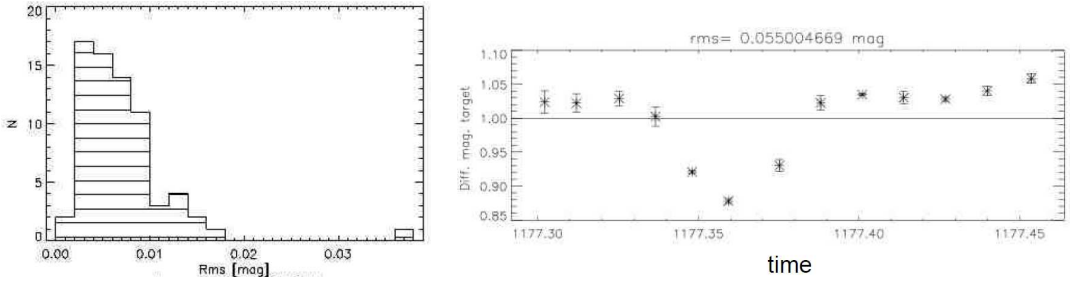
**Figure 5.** V mag distribution of the APACHE target sample.

mance: negligible temperature gradients, QE 80% in the whole wavelength range of interest. The open source observatory manager RTS2 (Kubanek 2010) is the choice for the high-level software control of the five-telescope system, including dynamic scheduling of the observations (for details, see Christille et al., this volume). Provisions have been made for the hardware of relevance for the APACHE operations control, data acquisition, processing centre and backup, which is necessary to accommodate the high data rate: two HP Z600 Workstation and one HP EliteBook 8740w Mobile Workstation as control PCs and workstations for data analysis, one HP Storage Works X1600 12TB as Network Attached Storage (NAS), connected together via Intranet 1 GB.

## 5 Survey Design: Input Catalogue and Observing Strategy

The APACHE Input Catalogue (AIC) of M dwarfs is constructed starting from the all-sky sample of 8889 bright ( $J < 10$ ) low-mass stars in Lépine & Gaidos (2011). The final AIC is composed of  $\sim 3000$  targets selected on the basis of their visibility from the OAVdA site and the availability of a suitable number of potentially good comparison stars in the telescopes' field-of-view. Moreover, cross-correlation with approximately two dozen catalogues was performed to get further information about the targets, such as a precise determination of spectral class, their projected rotational velocity  $V \sin i$ , their level of chromospheric activity, X-ray emission, the presence of binary companions detected either spectroscopically or by direct imaging, or astrometry, or a combination of these techniques. Based on accurate parameterization of Gaia's scanning law, the number of the Gaia transits scheduled was associated to each target. Based on all available information, a prioritized list was defined for the purpose of scheduling with the APACHE telescopes. The final optical magnitude distribution of the sample peaks at  $V \approx 13$  mag, with a median of  $V \sim 14$  mag (Fig. 5). The vast majority of targets is bright enough to allow for a host of additional characterization measurements, in case of positive detections of transiting planets. As a synergetic example, a carefully selected sub-sample of APACHE targets is currently the objective of high-precision RV monitoring with the HARPS-N spectrograph on the Telescopio Nazionale Galileo (TNG) as part of the long-term Global Architecture of Planetary Systems (GAPS) programme.

We built upon the experience matured with the pilot study to define and then refine the details of the observing strategy for the APACHE survey. Taking into account the use of synthetic transits, by assuming different numbers of consecutive exposures (from 1 to 4) and different temporal samplings



**Figure 6.** Left: rms photometric precision distribution for a sample of  $\sim 100$  APACHE targets observed during the first six months of survey operation. Right: APACHE#1 independent discovery lightcurve of an eclipsing binary first reported by Kiraga (2012).

(intervals of 20 through 50 minutes between two sets of consecutive pointings of the telescope on the same target), we evaluated the transit detection probabilities for stars with different average phase coverages in a given interval of possible orbital periods. The results of this analysis prompted us to adopt an initial observing strategy consisting of 3 exposures every  $\sim 20$  minutes, to be eventually revisited as the survey unfolds.

## 6 The First Six Months of Operations

Official survey operations started in July 2012, and the project evolution can be followed through its official website at <http://apacheproject.altervista.org>. With six months of data collected in hand, we can begin to look at some statistics of the survey. Approximately 100 fields have been observed thus far, with each APACHE telescope typically observing a dozen targets each night. The weather pattern of the Summer and Fall seasons has confirmed the expected fraction ( $\sim 50\%$ ) of useful nights.

Using the robust TEEPEE pipeline for the reduction and analysis of the photometric data (Giacobbe et al. 2012), we performed differential aperture photometry on the ensemble of targets observed to-date. The resulting distribution of the rms photometric precision for the sample (Fig. 6, left panel) has a median of  $\sim 5$  mmag, attesting of the quality of the data. We independently discovered (Fig. 6, right panel) a new eclipsing binary amongst our targets, that was published by Kiraga (2012) while we were performing a preliminary analysis on the data.

In parallel, we are developing the pipeline TSE (TEEPEE Sentinel) which is designed to perform real-time differential photometry of the targets during each night. Once fully developed and tested, TSE will be used for a prompt identification and assessment of interesting signals, triggering real-time changes in the observing strategy such as high-cadence follow-up of ongoing possible events (magnifying the sensitivity to long-period planets).

A significant update to the AIC is now ongoing, based on the latest spectroscopic information for many of the APACHE targets published by Lépine et al. (2013). The repercussions are particularly non-negligible in the M0-M1 sub-spectral type regime, with many objects being spectroscopically recognized as earlier-type dwarfs.

## 7 Summary

APACHE is the first Europe-based transit search for small-radius planets orbiting, bright, nearby M dwarfs. It draws on the pioneering work carried out by the MEarth project (Nutzman & Charbon-

neau 2008), in that it adopts a one target per field approach and the corresponding constraints on the observing strategy optimized for the array of five APACHE telescopes and the chosen site characteristics. It extends and complements MEarth in the Northern hemisphere, in that it targets early- to mid-M dwarfs, conceding in detectable planet radius for fixed photometric precision but focusing by construction on a target sample that's in principle more favourable for atmospheric characterization work in case of positive detection. After the first six months of operations, the program is delivering the expected quality in differential photometric precision. The synergy potential between APACHE, Doppler (e.g., HARPS-N), and astrometric (e.g., Gaia) techniques is potentially huge (see, e.g., Sozzetti, this volume), and the APACHE optical lightcurves of thousands of M dwarf stars will contribute directly to create a database of uniquely well-characterized low-mass stars in the solar neighborhood, inclusive of exoplanets with precise mass and radius determinations orbiting M dwarfs bright enough to enable detailed characterization of the planetary interior structure and atmospheres with upcoming space-borne observatories (e.g., CHEOPS, EChO, JWST).

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